TABLE 2. Brown Dwarf Candidates in the ρ Oph Cloud

Object	Q[a]	Sp. Ty. (±1.5)	log(T_{eff}) (±0.035)	(J-H) (mag)	(H-K) (mag)	K[b] (mag)	A_v[c] (mag)	log($L_{bol}/L_\odot$) (±0.16)	r_k[e]	Comments
CRBR 14	0.51	M7.5	3.400	1.75	1.20	12.28	10	-1.58	0.12	
CRBR 15	0.64	M5	3.467	2.46	1.95	11.94	17	-1.32	0.64	mm cont. source[f]
GY 5	0.55	M7	3.414	1.13	0.66	10.91	4.5	-1.20	-0.03	poss. x-ray source[g]
GY 10	0.47	M8.5	3.371	2.16	1.34	12.25	14	-1.44	-0.01	no mm cont. [h]
GY 11	0.58	M6.5	3.428	1.15	1.22	14.15	4.8	-2.69	0.63	H_2 emission lines[h]
GY 31	0.63	M5.5	3.454	···	4.00	13.10	56[d]	-0.15[d]	···	radio cont. source[i]
GY 37	0.59	M6	3.441	1.31	0.95	11.99	6.3	-1.62	0.18	poss. x-ray source[g]
GY 59	0.61	M6	3.441	1.86	1.21	11.68	11	-1.29	0.11	
GY 64	0.49	M8	3.386	1.87	1.43	13.33	11	-2.07	0.28	
GY 84	0.61	M6	3.441	2.40	1.26	11.63	16	-0.99	-0.13	
GY 107	0.76	M3	3.513	2.02	0.97	11.40	13	-0.87	-0.11	poss. background giant
GY 163	0.79	M2.5	3.524	···	2.28	12.66	32[d]	-0.75[d]	···	
GY 202	0.53	M7	3.414	2.09	1.70	12.97	13	-1.90	0.50	
GY 310	0.47	M8.5	3.371	1.29	0.83	11.08	5.7	-1.26	-0.01	
GY 326	0.71	M4	3.490	1.78	0.88	10.59	11	-0.70	-0.09	

[a]Water vapor index Q (see eqn. 1, Sec. 3.1).

[b]All photometry taken from Barsony *et al.* (1997) except that for GY31 taken from Comerón *et al.* (1993).

[c]Visual extinction $A_v = 9.09\times[(J-H)_{obs}$ - $(J-H)_0]$ except where noted.

[d]No J photometry is available so $A_v = 15.4\times[(H-K)_{obs}$ - $(H-K)_0]$, and L_{bol} is estimated from the absolute K magnitude plus the bolometric correction at K.

[e]Excess emission at λ=2.2 μm, $r_k = F_{Kex}/F_{K*}$, assuming no excess emission at H. Estimates are lower limits to the true values. See Sec. 3.2.2 for discussion.

[f]Millimeter continuum source with $S_{1.3mm}$=40 mJy (Motte, André, & Neri 1998).

[g]Candidate x-ray sources from Table 3 of Casanova *et al.* 1993.

[h]Subtraction of background emission reduces level of millimeter continuum emission at λ=1.3 mm to <10 mJy (Motte, André, & Neri 1998).

[i]Highly variable radio emission is observed at λ=6 cm by André (priv. comm.), probably arising from nonthermal geosynchrotron emission from a peristellar magnetosphere.